

THE ${}^7\text{Li}/{}^6\text{Li}$ ISOTOPE RATIO NEAR THE SUPERNOVA REMNANT IC 443¹

C. J. TAYLOR², A. M. RITCHHEY^{2,3}, S. R. FEDERMAN², AND D. L. LAMBERT⁴

Draft version March 1, 2013

ABSTRACT

We present an analysis of ${}^7\text{Li}/{}^6\text{Li}$ isotope ratios along four sight lines that probe diffuse molecular gas near the supernova remnant IC 443. Recent gamma-ray observations have revealed the presence of shock-accelerated cosmic rays interacting with the molecular cloud surrounding the remnant. Our results indicate that the ${}^7\text{Li}/{}^6\text{Li}$ ratio is lower in regions more strongly affected by these interactions, a sign of recent Li production by cosmic rays. We find that ${}^7\text{Li}/{}^6\text{Li} \approx 7$ toward HD 254755, which is located just outside the visible edge of IC 443, while ${}^7\text{Li}/{}^6\text{Li} \approx 3$ along the line of sight to HD 43582, which probes the interior region of the supernova remnant. No evidence of ${}^7\text{Li}$ synthesis by neutrino-induced spallation is found in material presumably contaminated by the ejecta of a core-collapse supernova. The lack of a neutrino signature in the ${}^7\text{Li}/{}^6\text{Li}$ ratios near IC 443 is consistent with recent models of Galactic chemical evolution, which suggest that the ν -process plays only a minor role in Li production.

Subject headings: ISM: abundances — ISM: atoms — ISM: individual objects (IC 443) — ISM: supernova remnants

1. INTRODUCTION

The production of Li involves contributions from a number of different astrophysical sources, including the Big Bang, Galactic cosmic rays (GCRs), red giant branch (RGB) and asymptotic giant branch (AGB) stars, and Type II supernovae (SNe II). The less abundant isotope, ${}^6\text{Li}$, is mainly a product of spallation (and $\alpha + \alpha$ fusion) reactions induced by GCRs in interstellar gas (e.g., Meneguzzi et al. 1971; Ramaty et al. 1997). Evidence for this can be seen in the solar system abundance ratio of ${}^6\text{Li}$ to ${}^9\text{Be}$, the sole stable isotope of Be which can be produced only through GCR spallation. The meteoritic ${}^6\text{Li}/{}^9\text{Be}$ ratio (5.6; Lodders 2003) is essentially the GCR value. In contrast, the ${}^7\text{Li}/{}^6\text{Li}$ ratio for solar system material is ~ 12 , while standard GCR nucleosynthesis predicts a ratio of ~ 1.5 . Since the primordial ${}^7\text{Li}$ abundance arising from the Big Bang is approximately a factor of 10 less than the present-day abundance, the majority of ${}^7\text{Li}$ must be produced in stars. However, the precise nature of the stellar source remains unclear.

The most promising candidates for a stellar ${}^7\text{Li}$ source include RGB and AGB stars, where ${}^7\text{Li}$ is produced via the Cameron-Fowler mechanism (e.g., Sackmann & Boothroyd 1999), and SNe II, where ${}^7\text{Li}$ is synthesized by neutrino-induced spallation in the He and C shells of the progenitor star during core collapse (i.e., the ν -process; Woosley et al. 1990). Direct observational evidence for the ν -process remains

elusive, yet could be revealed by detailed studies of ${}^7\text{Li}/{}^6\text{Li}$ isotope ratios in interstellar gas surrounding supernova remnants (SNRs). Since virtually no ${}^6\text{Li}$ is expected to be produced by neutrino nucleosynthesis (e.g., Woosley & Weaver 1995; Yoshida et al. 2008), interstellar material contaminated by SN II ejecta should exhibit a ${}^7\text{Li}/{}^6\text{Li}$ ratio that is enhanced over the ambient interstellar value. However, a reduced ${}^7\text{Li}/{}^6\text{Li}$ ratio might also be anticipated near a supernova remnant because SNRs are thought to be the primary sources responsible for cosmic-ray acceleration. Interactions between cosmic rays, accelerated by a supernova shock, and any nearby interstellar gas will drive the ${}^7\text{Li}/{}^6\text{Li}$ ratio toward the pure GCR value.

The well-studied SNR IC 443 provides an excellent opportunity to test which of these processes dominates the production of Li by SNe II. Located at a distance of 1.5 kpc in the Gem OB1 association, IC 443 is an intermediate-age (10–30 kyr) core-collapse SNR known to be interacting with atomic and molecular gas in its vicinity. Evidence of the interaction includes observations of shocked H I filaments (e.g., Braun & Strom 1986; Lee et al. 2008) and shocked molecular clumps (e.g., Huang et al. 1986; Dickman et al. 1992; Snell et al. 2005), as well as the detection of OH (1720 MHz) masers coincident with the shocked molecular material (Claussen et al. 1997; Hewitt et al. 2006). Many recent investigations of IC 443 have focused on the emission of gamma rays at GeV and TeV energies (e.g., Acciari et al. 2009; Tavani et al. 2010; Abdo et al. 2010). These studies strongly suggest a pionic origin for the gamma-ray emission, indicating that shock-accelerated cosmic rays (in addition to the shocks themselves) are interacting with the ambient molecular cloud.

In this Letter, we present an analysis of ${}^7\text{Li}/{}^6\text{Li}$ isotope ratios along four sight lines through IC 443. The ratios are extracted from high-resolution observations of Li I $\lambda 6707$ toward HD 254477, HD 254577, HD 43582,

¹ Based on observations obtained with the Hobby-Eberly Telescope, which is a joint project of the University of Texas at Austin, the Pennsylvania State University, Stanford University, Ludwig-Maximilians-Universität München, and Georg-August-Universität Göttingen.

² Department of Physics and Astronomy, University of Toledo, Toledo, OH 43606, USA; corbin.taylor@rockets.utoledo.edu; steven.federman@utoledo.edu

³ Department of Astronomy, University of Washington, Seattle, WA 98195, USA; aritchey@astro.washington.edu

⁴ W. J. McDonald Observatory, University of Texas at Austin, TX 78712, USA; dll@astro.as.utexas.edu

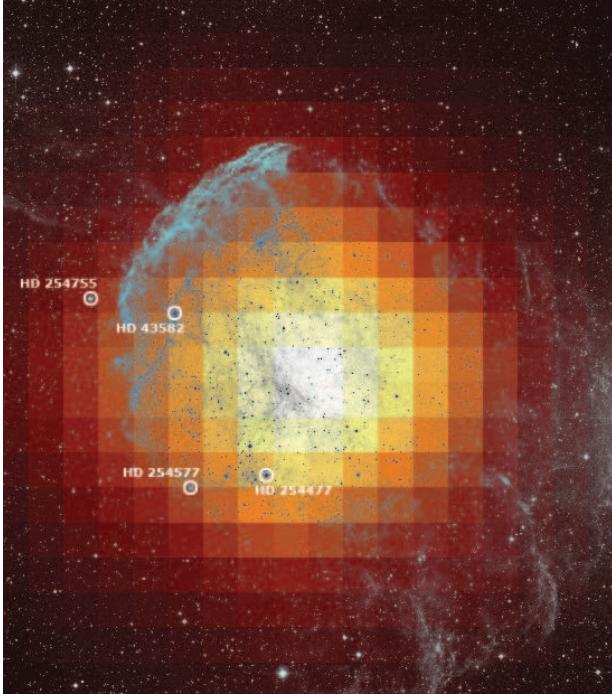


Figure 1. Composite image of IC 443. The optical image from the Digitized Sky Survey (POSS-II/red filter) is shown superimposed onto a gamma-ray intensity map in the 400 MeV to 3 GeV energy range obtained by the *Fermi* LAT (J. Hewitt 2011, private communication). The stars targeted for HET observations are labeled.

and HD 254755. Figure 1 shows the positions of the target stars in relation to the gamma-ray emission detected by the *Fermi* Large Area Telescope (LAT) in the 400 MeV to 3 GeV energy range (J. Hewitt 2011, private communication). Previous studies of the optical absorption profiles toward the selected targets have revealed extensive high-velocity gas (Welsh & Sallmen 2003; Hirschauer et al. 2009), indicating that these stars probe material directly associated with the SNR.

2. OBSERVATIONS AND DATA REDUCTION

The target stars were observed with the 9.2 m Hobby-Eberly Telescope (HET) at McDonald Observatory during the winters of 2008/2009 and 2009/2010. All observations employed the High Resolution Spectrograph (Tull 1998) with an effective slit width of $125 \mu\text{m}$ ($R \approx 120,000$) so that the fine structure lines of Li I could be adequately resolved. Two spectrographic settings (centered at 4931 \AA and 5936 \AA) provided data on Li I $\lambda 6707$. The setting at shorter wavelengths was needed to simultaneously obtain data on K I $\lambda 4044$, while the longer wavelength setting allowed the K I $\lambda 7698$ line to be observed. Both settings also yielded information on Ca I $\lambda 4226$, CH⁺ $\lambda 4232$, and CH $\lambda 4300$.

Basic information concerning the target stars can be found in Hirschauer et al. (2009). The three brighter targets were observed with both instrumental setups, resulting in total exposure times on Li I $\lambda 6707$ of 6.1 hr for HD 254577 and 4.7 hr for HD 43582 and HD 254755. These exposure times yielded signal-to-noise ratios (S/N) per resolution element near Li I of about 1100. The faintest star in our sample, HD 254477, was observed only with the longer wavelength setting for a total of 2.0

hr, resulting in a S/N of 380.

Standard procedures within IRAF were employed for bias correction, cosmic-ray removal, scattered light subtraction, and flat fielding. Ideally, a master flat was obtained from the median of flats taken for a given night. In some instances, only the first flat of a sequence could be used due to the appearance of an emission feature near 6708 \AA in subsequent flats. One-dimensional spectra were extracted from the processed images and were wavelength-calibrated after identifying emission lines in Th-Ar comparison spectra. Before shifting the calibrated spectra to the local standard of rest (LSR) frame, small velocity corrections were applied based on the measured wavelengths of the atmospheric [O I] emission lines at 6300 \AA and 5577 \AA . This was necessary to account for offsets in the velocity zero points caused by slight changes in spectrograph placement between the stellar and Th-Ar exposures.

Final spectra were produced by co-adding the individual exposures of a given target for orders containing the lines of interest. The co-added spectra were normalized to unity by fitting low-order polynomials to regions free of interstellar or telluric features. Figure 2 presents the reduced Li I spectra for the four sight lines probing IC 443. From the widths of Th I emission lines in the nightly comparison spectra, we found that the actual resolving power of the spectrograph during our observations was $R = 98,000$.

3. ANALYSIS

Detailed knowledge of the velocity component structure in each direction is necessary if meaningful $^7\text{Li}/^6\text{Li}$ ratios are to be extracted from the complicated Li I line profiles, which exhibit fine and hyperfine structure in addition to isotopic splitting. For a single, optically-thin cloud, and in the absence of ^6Li , the two fine-structure components of the Li I doublet will exhibit relative strengths of 2:1 and a separation of 0.15 \AA . Addition of ^6Li reduces this ratio because the stronger of the two ^6Li components is superimposed onto the weaker of the components from ^7Li (i.e., the isotope shift is approximately equal to the fine-structure separation).

Further complications arise when multiple clouds are present along the line of sight. Thus, the first step in our analysis was to derive a robust solution for the line-of-sight component structure using species with moderately strong absorption lines. The K I $\lambda 7698$ and CH $\lambda 4300$ lines were of particular interest as these species are expected to coexist with Li I in cool, diffuse clouds (Welty & Hobbs 2001; Knauth et al. 2003; Pan et al. 2005). While the weaker K I line at 4044 \AA would provide a better template for Li I $\lambda 6707$ (due to these lines having similar intrinsic strengths), absorption from K I $\lambda 4044$ is detected only toward HD 254755. We also analyzed the Ca I $\lambda 4226$ and CH⁺ $\lambda 4232$ absorption profiles to check for consistency in velocity among the components in common with K I and CH.

The column densities and component parameters for all of the relevant species were determined through profile synthesis using the program ISMOD developed by Y. Sheffer (see, e.g., Sheffer et al. 2008). ISMOD treats the velocities, b -values, and column densities of the absorption components as free parameters while minimizing the rms deviations in the residuals of the fit. Hyperfine

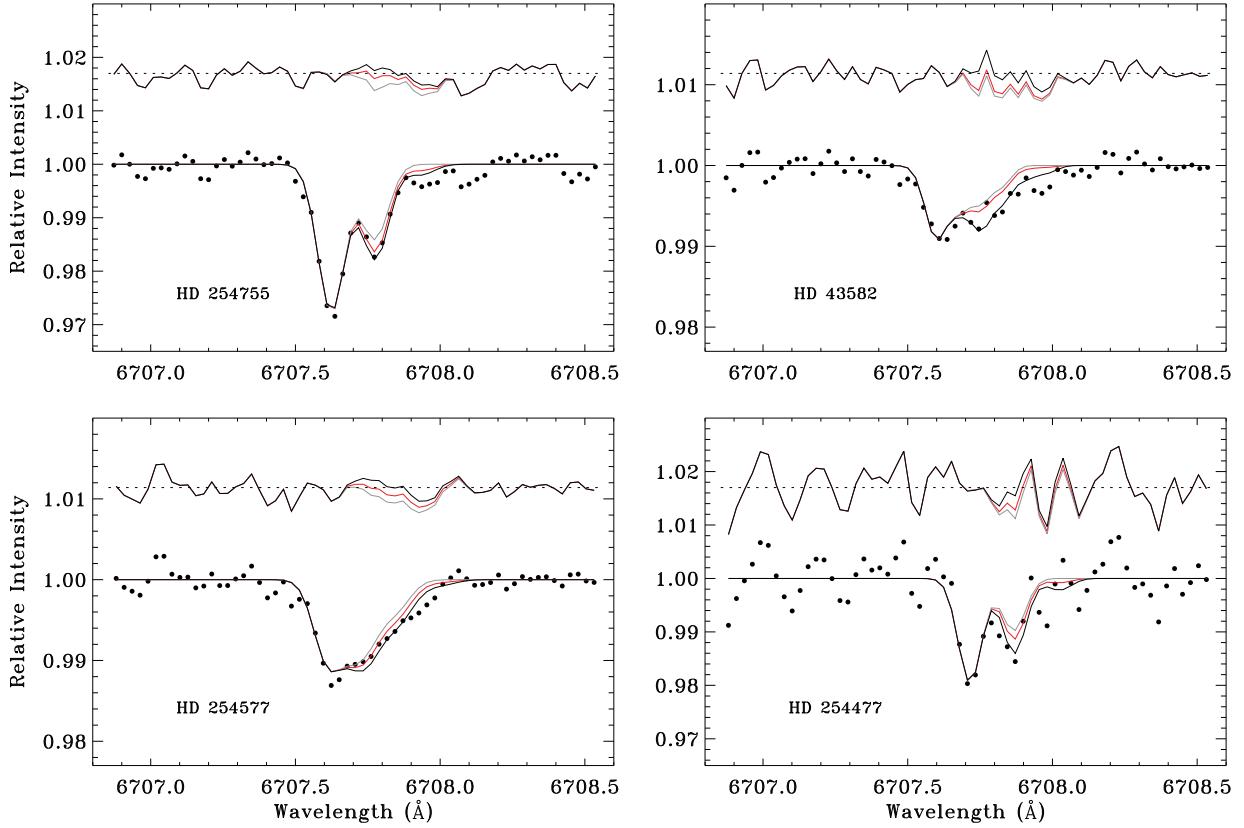


Figure 2. Profile synthesis fits to the fine structure lines of ${}^7\text{Li}$ I and ${}^6\text{Li}$ I toward HD 254755, HD 43582, HD 254577, and HD 254477. Along with the best-fit synthetic profile (black line), two alternative syntheses are shown in each case, one assuming a solar system ${}^7\text{Li}/{}^6\text{Li}$ ratio (red line) and one assuming no ${}^6\text{Li}$ is present (grey line). The parameters for ${}^7\text{Li}$ in these alternative syntheses are the same as in the best-fit cases. Residuals are given above each fit. For HD 43582 and HD 254577, the best-fit profile displayed is that based on the K I template. Note that the alternative syntheses are less satisfactory in general, but particularly for HD 43582.

structure is included in our synthetic profiles for Li I and K I, with wavelengths and f -values adopted from Morton (2003). Our fits also account for Λ -doubling in the CH line. To extract the ${}^7\text{Li}/{}^6\text{Li}$ ratio, the profile synthesis routine performs a simultaneous fit to the blended fine structure lines of ${}^7\text{Li}$ I and ${}^6\text{Li}$ I. The component structure is assumed to be identical for the two isotopes, with the isotope ratio left as a free parameter. Our results for K I and CH were used to determine the number of velocity components to include in the synthesis for Li I.

For HD 254755, a single component contributes more than 90% of the total column density of K I and CH. Our best one-component fit to the Li I profile (Figure 2) yields a ${}^7\text{Li}/{}^6\text{Li}$ ratio of 7.1 ± 2.4 , where the errors are dominated by the observational uncertainties associated with the detection of ${}^6\text{Li}$. The velocity and b -value resulting from the fit (see Table 1) are in good agreement with those for the corresponding component in K I ($v_{\text{LSR}} = -6.4 \text{ km s}^{-1}$; $b = 1.2 \text{ km s}^{-1}$) and CH ($v_{\text{LSR}} = -5.9 \text{ km s}^{-1}$; $b = 2.1 \text{ km s}^{-1}$).

For HD 43582 and HD 254577, two velocity components comprise 70–80% of the total columns of K I and CH. The Li I profiles in these directions are manifestly more blended than toward HD 254755, indicating that multiple components are present. We therefore kept the component structure fixed in the Li I syntheses, with component parameters determined separately from the two dominant components in K I and CH. The results are given in Table 1 and details concerning the fits are pre-

sented in Figures 3 and 4 for HD 43582 and HD 254577, respectively. In both cases, the ${}^7\text{Li}/{}^6\text{Li}$ ratios derived from the K I and CH templates are indistinguishable given the uncertainties. Thus, we adopt the average ratios for further analysis (i.e., ${}^7\text{Li}/{}^6\text{Li} = 3.1 \pm 1.4$ for HD 43582 and 6.1 ± 3.2 for HD 254577).

For HD 254477, the strongest velocity component accounts for only about 60% of the K I and CH column density, but no other component makes up more than 10% of the profile in these species. Our best one-component fit to the Li I profile yields a consistent velocity and b -value compared to the dominant component in K I ($v_{\text{LSR}} = -2.8 \text{ km s}^{-1}$; $b = 2.3 \text{ km s}^{-1}$) and CH ($v_{\text{LSR}} = -2.5 \text{ km s}^{-1}$; $b = 1.0 \text{ km s}^{-1}$). However, the significantly lower S/N makes it impossible to confidently derive a ${}^7\text{Li}/{}^6\text{Li}$ ratio in this direction. Nominally, we find ${}^7\text{Li}/{}^6\text{Li} = 4.5 \pm 3.6$, but with $N({}^6\text{Li I}) = (2.2 \pm 1.8) \times 10^9 \text{ cm}^{-2}$ (i.e., not even a 2σ detection). A more appropriate 3σ upper limit on $N({}^6\text{Li I})$ of $\lesssim 5.4 \times 10^9 \text{ cm}^{-2}$ gives a lower limit on ${}^7\text{Li}/{}^6\text{Li}$ of $\gtrsim 1.9$.

4. DISCUSSION AND CONCLUSIONS

Our results on ${}^7\text{Li}/{}^6\text{Li}$ isotope ratios in the vicinity of IC 443 shed new light on Li production associated with SNe II. Before these observations, it was not clear whether one should expect an enhanced abundance of ${}^7\text{Li}$ near SNe II, resulting from neutrino spallation during the core-collapse phase, or a higher relative abundance

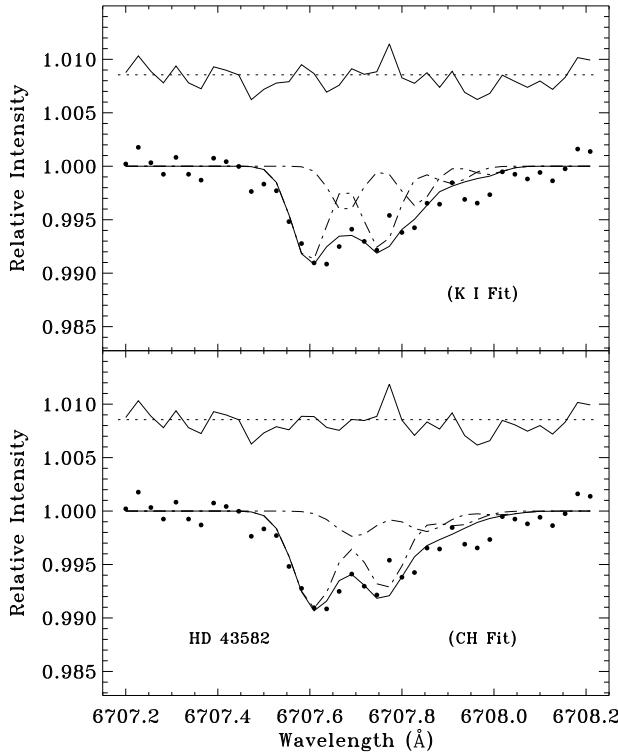


Figure 3. Synthesis fits to the Li I profile toward HD 43582, based on the K I template (upper panel) and the CH template (lower panel). The dot-dashed lines show the contributions from individual cloud components to the overall profile, shown as a solid line.

of ${}^6\text{Li}$, due to spallation and fusion reactions initiated by shock-accelerated cosmic rays. Our analysis of Li I absorption along four sight lines through IC 443 suggests that cosmic-ray interactions dominate the production of Li at least for this particular SNR.

All of the ${}^7\text{Li}/{}^6\text{Li}$ ratios we obtain near IC 443 are lower than the solar system (meteoritic) value of 12.2 (Lodders 2003), which Knauth et al. (2003) suggest is representative of gas in the solar neighborhood. While our results for HD 254577 and HD 254755 are actually consistent with the solar system ratio at the 2σ level, the ratio we find toward HD 43582 is lower by more than 6σ . There is evidence, however, that the Li isotope ratio in the interstellar medium (ISM) in general is lower than the solar system value. The weighted mean ISM ratio from measurements available in the literature (Lemoine et al. 1993; Meyer et al. 1993; Lemoine et al. 1995; Knauth et al. 2000; Howarth et al. 2002; Knauth et al. 2003; Kawanomoto et al. 2009) is 7.3 ± 0.6 (if one considers only the strongest absorption component in each case). This value is very similar to the ratio we find toward HD 254755, which lies beyond the outer edge of the H α emission contours in the northeast of IC 443 (see Figure 1). The gas in this direction may thus represent material relatively unaffected by the supernova shock, implying that a ${}^7\text{Li}/{}^6\text{Li}$ ratio of ~ 7 characterized the ambient molecular cloud before the occurrence of the supernova.

In contrast, the ${}^7\text{Li}/{}^6\text{Li}$ ratio toward HD 43582 is 40–80% lower than the mean value observed in the local ISM. The stars HD 254755 and HD 43582 are separated

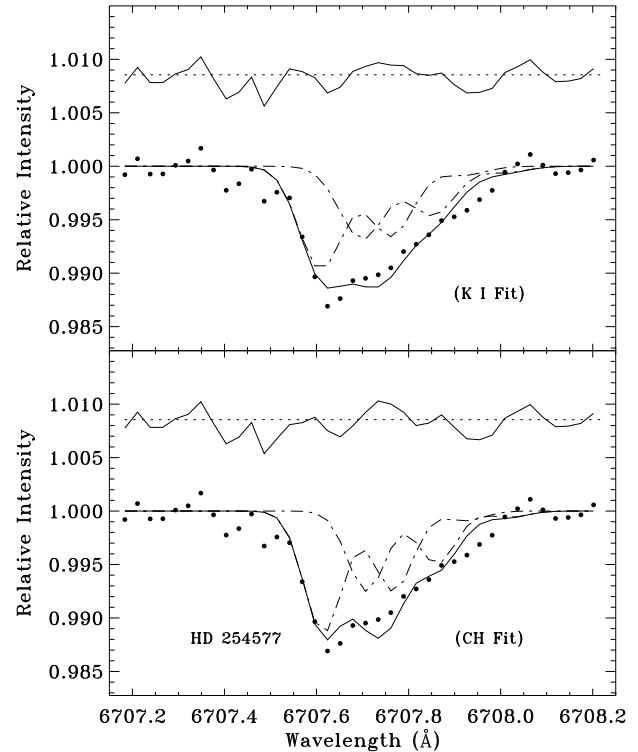


Figure 4. Same as Figure 3 except for the Li I profile toward HD 254577.

by only $7'.4$ on the sky (or 3.2 pc at the distance of IC 443), but the line of sight to HD 43582 clearly penetrates the interior region of the SNR. In fact, HD 43582 is the only star of those investigated by Hirschauer et al. (2009) that exhibits Ca II absorption at high positive velocity, strongly suggesting that it lies behind the receding edge of the expanding shell of the remnant. If the molecular gas in this direction has been subjected to a more intense cosmic-ray irradiation than that toward HD 254755, then an enhancement in the relative abundance of ${}^6\text{Li}$ would be a natural consequence. Indeed, the ${}^7\text{Li}/{}^6\text{Li}$ ratio toward HD 43582 approaches the value predicted by models of GCR nucleosynthesis (~ 1.5 ; Meneguzzi et al. 1971; Ramaty et al. 1997; Lemoine et al. 1998). The sight lines to HD 254577 and HD 254477 presumably also pass through the interior of the SNR. However, given the uncertainties, the ${}^7\text{Li}/{}^6\text{Li}$ ratios in these directions are consistent with both the mean ISM value (as observed toward HD 254755) and the lower value we find toward HD 43582.

The discovery of a low ${}^7\text{Li}/{}^6\text{Li}$ ratio in IC 443 corroborates the conclusions drawn from gamma-ray emission studies of the region (e.g., Acciari et al. 2009; Tavani et al. 2010; Abdo et al. 2010), which strongly suggest that the gamma radiation results from the decay of neutral pions produced through cosmic-ray interactions with molecular gas. Complementary results on the cosmic-ray ionization rate were recently reported by Indriolo et al. (2010). These authors find large H $_3^+$ column densities near IC 443 and deduce an H $_2$ ionization rate of $\zeta_2 \approx 2 \times 10^{-15} \text{ s}^{-1}$, or about five times the rate typically found for diffuse molecular clouds (e.g., Indriolo et al. 2007).

Both the H $_3^+$ abundances discussed by Indriolo et al.

(2010) and the ${}^7\text{Li}/{}^6\text{Li}$ ratios presented here indicate regions of enhanced cosmic-ray activity within IC 443. However, the two observables present potentially interesting differences in terms of which sight lines exhibit the largest enhancements. Indriolo et al. (2010) infer an ionization rate of $\zeta_2 = 2.6^{+1.3}_{-1.9} \times 10^{-15} \text{ s}^{-1}$ (the highest in their survey) toward HD 254577, yet report only an upper limit on ζ_2 toward HD 43582 since H_3^+ is not detected. (The upper limit is $\lesssim 1.4 \times 10^{-15} \text{ s}^{-1}$ if uncertainties in the H_2 column density are included.) These results appear to be in conflict with our determinations of ${}^7\text{Li}/{}^6\text{Li}$ for these directions, since we find a lower ${}^7\text{Li}/{}^6\text{Li}$ ratio toward HD 43582 than toward HD 254577. However, the uncertainties both in our determinations and in those of Indriolo et al., which rely on a number of assumptions concerning the physical conditions in the clouds containing H_3^+ , are large. Moreover, since the abundance of H_3^+ traces the current cosmic-ray ionization rate, while the Li isotope ratio probes the integrated cosmic-ray flux, some differences in these two observables might be expected (B. Fields 2011, private communication).

A decrease in the ${}^7\text{Li}/{}^6\text{Li}$ ratio due to enhanced cosmic-ray activity (resulting in newly-synthesized ${}^6\text{Li}$ and ${}^7\text{Li}$) should be accompanied by an increase in the elemental Li abundance. If the clouds toward HD 43582 initially possessed a ${}^7\text{Li}/{}^6\text{Li}$ ratio of ~ 7 , then the presently-observed ratio of 3.1 would imply a factor-of-two increase in the elemental Li abundance. Some evidence for this effect can be seen in the $N(\text{Li I})/N(\text{K I})$ ratios found toward the stars in our sample: $(2.7 \pm 0.2) \times 10^{-3}$ for HD 254755, $(5.9 \pm 0.6) \times 10^{-3}$ for HD 254577, $(6.4 \pm 0.9) \times 10^{-3}$ for HD 43582, and $(6.8 \pm 1.4) \times 10^{-3}$ for HD 254477. The three sight lines that pass through the interior of the remnant seem to exhibit the expected factor-of-two enhancement in the abundance of Li (relative to K) if the initial abundances were similar to that seen toward HD 254755.

Still, none of the $N(\text{Li I})/N(\text{K I})$ ratios we find near IC 443, are entirely unusual compared to values typical of diffuse clouds (e.g., Welty & Hobbs 2001; Knauth et al. 2003). Without directly determining the elemental Li/K ratios, which would require making uncertain corrections for ionization and depletion onto grains, it is difficult to conclude whether an enhancement in Li is indeed observed. We can, however, state that a localized enhancement would be energetically feasible. If we assume that $n_{\text{H}} \sim 200 \text{ cm}^{-3}$ for the gas toward HD 43582 (as adopted by Indriolo et al. 2010), then a factor-of-two increase in the Li abundance would require that $\sim 3 \times 10^{50} \text{ erg}$ are imparted to the accelerated cosmic rays by the SN shock (see Ramaty et al. 2000). This is only $\sim 20\%$ of the typical mechanical energy available in SN II ejecta (Woosley & Weaver 1995)

Our observations suggest that the abundance of ${}^6\text{Li}$ relative to ${}^7\text{Li}$ has been enhanced near IC 443 by interactions between shock-accelerated cosmic rays and the ambient molecular cloud. At the same time, we find no evidence of ${}^7\text{Li}$ synthesis by neutrino-induced spallation in material that presumably has been contaminated by the ejecta of a core-collapse supernova. Given the age of the remnant, the hot ejecta should have had ample time to cool, as it interacts with its dense surroundings, to the point at which absorption lines from neutral atoms would become visible. For comparison,

Wallerstein et al. (1995) estimated a cooling time of 800 yr for the Vela remnant, which is about the same age as IC 443 but where the ambient ISM is of much lower density. The lack of a neutrino signature is consistent with recent models of Galactic chemical evolution (e.g., Prantzos 2012), which suggest a very minor role for the ν -process in producing ${}^7\text{Li}$. Future measurements of ${}^7\text{Li}/{}^6\text{Li}$ in gas surrounding other SNRs will help to establish unequivocally the role that neutrino spallation plays in Li production.

We gratefully acknowledge useful conversations with Brian Fields and Nick Indriolo. We also thank Jack Hewitt for providing us with the *Fermi* data for Figure 1. C. Taylor participated in the Research Experience for Undergraduates program at the University of Toledo under NSF-REU grant PHY-1004649. D.L.L. acknowledges support from the Robert A. Welch Foundation of Houston, Texas through grant F-634. The Hobby-Eberly Telescope is a joint project of the University of Texas at Austin, the Pennsylvania State University, Stanford University, Ludwig-Maximilians-Universität München, and Georg-August-Universität Göttingen. The HET is named in honor of its principal benefactors, William P. Hobby and Robert E. Eberly.

REFERENCES

- Abdo, A. A., et al. 2010, ApJ, 712, 459
- Acciari, V. A., et al. 2009, ApJ, 698, L133
- Braun, R., & Strom, R. G. 1986, A&A, 164, 193
- Claussen, M. J., Frail, D. A., Goss, W. M., & Gaume, R. A. 1997, ApJ, 489, 143
- Dickman, R. L., Snell, R. L., Ziurys, L. M., & Huang, Y.-L. 1992, ApJ, 400, 203
- Hewitt, J. W., Yusef-Zadeh, F., Wardle, M., Roberts, D. A., & Kassim, N. E. 2006, ApJ, 652, 1288
- Hirschauer, A., Federman, S. R., Wallerstein, G., & Means, T. 2009, ApJ, 696, 1533
- Howarth, I. D., Price, R. J., Crawford, I. A., & Hawkins, I. 2002, MNRAS, 335, 267
- Huang, Y.-L., Dickman, R. L., & Snell, R. L. 1986, ApJ, 302, L63
- Indriolo, N., Geballe, T. R., Oka, T., & McCall, B. J. 2007, ApJ, 671, 1736
- Indriolo, N., Blake, G. A., Goto, M., Usuda, T., Oka, T., Geballe, T. R., Fields, B. D., & McCall, B. J. 2010, ApJ, 724, 1357
- Kawanomoto, S., et al. 2009, ApJ, 701, 1506
- Knauth, D. C., Federman, S. R., & Lambert, D. L. 2003, ApJ, 586, 268
- Knauth, D. C., Federman, S. R., Lambert, D. L., & Crane, P. 2000, Nature, 405, 656
- Lee, J.-J., Koo, B.-C., Yun, M. S., Stanimirović, S., Heiles, C., & Heyer, M. 2008, AJ, 135, 796
- Lemoine, M., Ferlet, R., & Vidal-Madjar, A. 1995, A&A, 298, 879
- Lemoine, M., Ferlet, R., Vidal-Madjar, A., Emerich, C., & Bertin, P. 1993, A&A, 269, 469
- Lemoine, M., Vangioni-Flam, E., & Cassé, M. 1998, ApJ, 499, 735
- Lodders, K. 2003, ApJ, 591, 1220
- Meneguzzi, M., Audouze, J., & Reeves, H. 1971, A&A, 15, 337
- Meyer, D. M., Hawkins, I., & Wright, E. L. 1993, ApJ, 409, L61
- Morton, D. C. 2003, ApJS, 149, 205
- Pan, K., Federman, S. R., Sheffer, Y., & Andersson, B.-G. 2005, ApJ, 633, 986
- Prantzos, N. 2012, A&A, submitted
- Ramaty, R., Kozlovsky, B., Lingenfelter, R. E., & Reeves, H. 1997, ApJ, 488, 730
- Ramaty, R., Scully, S. T., Lingenfelter, R. E., & Kozlovsky, B. 2000, ApJ, 534, 747
- Sackmann, I.-J., & Boothroyd, A. I. 1999, ApJ, 510, 217
- Sheffer, Y., Rogers, M., Federman, S. R., Abel, N. P., Gredel, R., Lambert, D. L., & Shaw, G. 2008, ApJ, 687, 1075
- Snell, R. L., Hollenbach, D., Howe, J. E., Neufeld, D. A., Kaufman, M. J., Melnick, G. J., Bergin, E. A., & Wang, Z. 2005, ApJ, 620, 758
- Tavani, M., et al. 2010, ApJ, 710, L151
- Tull, R. G. 1998, Proc. SPIE, 3355, 387

- Wallerstein, G., Vanture, A. D., Jenkins, E. B., & Fuller, G. M. 1995, ApJ, 449, 688
- Welsh, B. Y., & Sallmen, S. 2003, A&A, 408, 545
- Welty, D. E., & Hobbs, L. M. 2001, ApJS, 133, 345
- Woosley, S. E., Hartmann, D. H., Hoffman, R. D., & Haxton, W. C. 1990, ApJ, 356, 272
- Woosley, S. E., & Weaver, T. A. 1995, ApJS, 101, 181
- Yoshida, T., Suzuki, T., Chiba, S., Kajino, T., Yokomakura, H., Kimura, K., Takamura, A., & Hartmann, D. H. 2008, ApJ, 686, 448

Table 1
Li I Profile Synthesis Results

Star	Template	v_{LSR} (km s $^{-1}$)	b (km s $^{-1}$)	$N(^7\text{Li I})$ (10 9 cm $^{-2}$)	${}^7\text{Li}/{}^6\text{Li}$	$N(\text{Li I})$ (10 9 cm $^{-2}$)
HD 254477	...	-2.2	1.7	10.2 ± 1.8	$\gtrsim 1.9$	$12.4 \pm 2.5^{\text{a}}$
HD 254577	K I	(-6.9) (-2.9)	(2.3) (2.4)	5.8 ± 0.7 4.2 ± 0.7	5.7 ± 3.1	11.8 ± 1.3
	CH	(-6.7) (-2.5)	(1.8) (1.7)	6.1 ± 0.6 3.9 ± 0.6	6.4 ± 3.4	11.6 ± 1.1
HD 43582	K I	(-7.4) (-3.9)	(1.4) (0.6)	4.3 ± 0.6 1.8 ± 0.5	2.8 ± 1.0	8.3 ± 1.1
	CH	(-6.9) (-2.9)	(2.0) (1.8)	5.1 ± 0.7 1.3 ± 0.6	3.3 ± 1.7	8.3 ± 1.3
HD 254755	...	-6.3	2.2	16.6 ± 0.8	7.1 ± 2.4	18.9 ± 1.1

Note. — Velocities and b -values shown in parentheses were held fixed during profile synthesis. The last column gives the total Li I column density. The corresponding K I column densities for components also seen in Li I are 18.3 ± 0.1 , 19.7 ± 0.1 , 12.9 ± 0.1 , and 69.4 ± 1.7 (in units of 10^{11} cm $^{-2}$) for HD 254477, HD 254577, HD 43582, and HD 254755, respectively.

^a Assumes ${}^7\text{Li}/{}^6\text{Li} \sim 4.5$.